

GALACTIC S STARS: INVESTIGATIONS OF COLOR, MOTION, AND SPECTRAL FEATURES

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ABSTRACT

Known bright S stars, recognized as such by their enhanced s-process abundances and C/O ratio, are typically members of the asymptotic giant branch (AGB) or the red giant branch (RGB). Few modern digital spectra for these objects have been published, from which intermediate resolution spectral indices and classifications could be derived. For published S stars we find accurate positions using the Two-Micron All Sky Survey (2MASS), and use the FAST spectrograph of the Tillinghast reflector on Mt. Hopkins to obtain the spectra of 57 objects. We make available a digital S star spectral atlas consisting of 14 spectra of S stars with diverse spectral features. We define and derive basic spectral indices that can help distinguish S stars from late-type (M) giants and carbon stars. We convolve all our spectra with the Sloan Digital Sky Survey (SDSS) bandpasses, and employ the resulting *gri* magnitudes together with 2MASS *JHK_s* mags to investigate S star colors. These objects have colors similar to carbon and M stars, and are therefore difficult to distinguish by color alone. Using near and mid-infrared colors from IRAS and AKARI, we identify some of the stars as intrinsic (AGB) or extrinsic (with abundances enhanced by past mass-transfer). We also use *V* band and 2MASS magnitudes to calculate a temperature index for stars in the sample. We analyze the proper motions and parallaxes of our sample stars to determine upper and lower limit absolute magnitudes and distances, and confirm that most are probably giants.

Subject headings: Stars: abundances, AGB and post-AGB, carbon

1. INTRODUCTION

When a low to intermediate mass star exhausts its core supply of hydrogen, contraction and heating of the core causes the outer layers of the star to expand and cool, whereby the star becomes a red giant (RG) with an average effective temperature of 5000 K or lower. Once the core temperature rises to about 3×10^8 K, helium burning begins, and the stellar surface contracts and heats creating the horizontal branch (for Population II) or red clump (for Population I) stars on the color-magnitude diagram. Completion of core helium burning and the start of helium shell burning returns the star to a luminous, cooler phase as it evolves onto the asymptotic giant branch (early- or E-AGB) where the star greatly expands and the hydrogen shell is (almost) extinguished. Later, as the He-burning shell approaches the H-He discontinuity, a phase of double shell burning begins. Recurrent thermal instabilities in this thermally-pulsing asymptotic giant branch (TP-AGB) phase are driven by helium shell flashes, periodic thermonuclear runaway events in the He-shell (Iben & Renzini 1983). The energy released by these pulses expand the star, whereby the hydrogen shell is again basically extinguished during some time. A third dredge up might then occur, and afterward the hydrogen shell will re-ignite (and a new thermal pulse will occur). During these short-lived pulses, nucleosynthesis products from combined H-shell and He-shell burning are dredged up to the outer layers of the star. The complex details of the dredge-up and its outcome are strongly mass-dependent (Herwig 2005 and references therein), but generally speaking, surface abundance enhancements result in C, He, and the s-process elements, including Ba, La, Zr, and Y.

S and Carbon (C) stars are traditionally thought to be members of the TP-AGB. In addition to possibly elevated C/O ratios, these stars exhibit other unique spectral characteristics as a result of the dredge-up. While M giants often exhibit titanium oxide (TiO) absorption bands, both carbon and zirconium have a higher affinity to free oxygen. Therefore, a higher C/O ratio also results in the disappearance of TiO bands and the appearance of zirconium oxide (ZrO) bands. S stars are therefore distinguished primarily by their ZrO bands, while C stars show strong C₂ and CN bands. The S star spectral class is divided into three subtypes (in order of increasing carbon abundance): MS, S, and SC stars. Classification into one particular class is difficult, but is generally based on the C/O ratio and the relative strengths of TiO, ZrO, and CN molecular absorption bands. Overall, a star in the S star class is generally defined as having a C/O ratio (*r*) of $0.5 < r < 1.0$ and strong ZrO absorption features (Jorissen et al. 1993; Van Eck et al. 2010). Pure S stars are those that show only ZrO bands, and no TiO bands (Wyckoff & Clegg 1978). Classification is complicated by the fact that increased intensity of the ZrO bands may be due to an excess of Zr, rather than being exclusively tied to the C/O ratio (Plez et al. 2003). Furthermore, many AGB S and C giants are highly variable, and since the strength of molecular bands depends on the effective temperature and surface gravity of the star, spectral type can vary significantly with time.

S stars whose spectra show lines of the short-lived element s-process technetium (Tc) are known as "intrinsic" S stars, and are most likely be members of the TP-AGB. ⁹⁹Tc is only produced by the s-process, while the star is on the AGB. Since the half-life of ⁹⁹Tc is only 2.13×10^5 years, and the average duration of the AGB phase of stellar evolution is roughly 1 Myr, stars exhibiting significant abundances of technetium are almost certainly members of the AGB (Van Eck & Jorissen 2000). By contrast, technetium-poor S stars are known as "extrinsic" S stars. These stars are likely to be part of binary sys-

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tems and their unusual chemical abundances probably originated in a past mass-transfer episode (Van Eck & Jorissen 2000). The present extrinsic S star likely once accreted s -process rich material from its companion, a TP-AGB star at the time that has since evolved into a white dwarf. Extrinsic S stars therefore display enriched s -process elements, despite having never produced these elements themselves. Most known extrinsic S stars are assumed to be members of the red giant branch (RGB). They can sometimes be distinguished from intrinsic S stars on the basis of color, because they do not necessarily show the red excess characteristic of stars on the AGB. While all currently known extrinsic S stars are thought to be giants, these definitions open up the possibility of extrinsic S stars on the main sequence (hereafter referred to as dwarf S stars) that have previously accreted s -process material and carbon from a companion but whose Tc has since decayed. Analogous to the dwarf C stars (Green et al. 1991), a faint S star could be shown to be a dwarf if it has a sufficiently large parallax or proper motion; a large proper motion predicts a tangential velocity greater than the Galactic escape speed unless the object is faint and nearby.

S stars are very similar to C stars, with the major difference being a slightly lower C/O ratio. Faint high galactic latitude carbon (FHLC) stars were once assumed to be giant stars at large distances. In 1977, however, G77-61 was discovered to have high proper motion and an upper limit absolute magnitude of +9.6, therefore making it the first known carbon dwarf star (Dahn et al. 1977). In this case, the term dwarf refers to a star on the main sequence. Green et al. (1991) discovered four other dwarf carbon (dC) stars using techniques that selected known carbon stars for high proper motion. Now, well over 100 dwarf carbon stars are known (Downes et al. 2004). The local space density of dCs is far higher than all types of C giants combined (Green et al. 1992). Therefore, contrary to previous assumptions, dwarf C stars are the numerically dominant type of carbon star in the Galaxy.

Since dC stars are now known to be so common, we have begun a search for dwarf S stars. There is little reason to believe that dwarf S stars should not exist. In fact, they may be more abundant than dC stars, since less carbon enhancement is required to produce an S dwarf as compared to a C dwarf. On the other hand, if the range of abundance ratios that produces the characteristic features of S stars is narrow, they may be quite rare. Furthermore, the range of C/O ratios that produces S star spectral features may be different for giants and dwarfs, given the higher gravities and larger temperature ranges in dwarfs.

We use medium-resolution digital spectra of known S stars (which are probably giants) to generate corresponding Sloan Digital Sky Survey (SDSS) colors, to see whether their colors might be sufficiently distinctive to search for additional S stars in the SDSS with reasonable efficiency. We also use the FAST spectra to investigate molecular band spectral indices for classification.

Our paper is organized as follows: in §2 we briefly review the observations and our data reduction and processing procedures. In section §3 we describe the production of a spectral atlas of S stars, and derive spectral indices from the molecular band strengths in §4. In §5 we describe the generation of SDSS colors and analyze the overall colors of the FAST sample in order to determine whether an efficient color selection for S stars exist. In §6 we discuss techniques for distinguishing intrinsic and extrinsic S stars and apply them to the FAST sample. In §8 we analyze the parallaxes and proper motions

of S stars in the sample to characterize upper and lower limit magnitudes and distances. We use colors to generate approximate temperature indices for S stars in §7. In §9 we summarize our results and present our conclusions.

2. OBSERVATIONS

2.1. Sample Selection

The largest spectral catalog of S stars is Stephenson's second edition (1984), which includes 1347 objects with poorly known positions. The stars in this catalog were originally identified from blue, red, and infrared objective prism plates. The majority of the S stars were discovered on the basis of the red system of ZrO absorption bands, with a bandhead around 6474 Å. Since the Stephenson catalog has a limiting V magnitude of approximately 11.5, most of the objects are highly saturated in the SDSS and reliable u , g , r , i , and z magnitudes have not been found for any known S stars. We assembled a sample of known S giants in order to generate likely g and r colors for S giants and dwarfs. These constraints on colors of S giants and dwarfs could allow us to search the SDSS and 2MASS catalogs for potential dwarf S star candidates.

To choose the stars making up the sample, we selected high Galactic latitude ($|b| > 20$ deg), northern ($\delta > -05$ deg) S stars from Stephenson's catalog. This ensures that our sample is accessible from Mt. Hopkins, and also decreases problems of object confusion or reddening. We then correlated the sample from Stephenson's catalog with 2MASS to find improved positions for most objects; accuracy of the original positions is typically $\sim 4''$, but commonly as poor as $20''$. The final sample list, which includes 57 objects, can be found in Table 1. Each object is accompanied by its identification in both the Stephenson and 2MASS catalogs, as well as its position. Since many of these stars are known to be variable, we include the date of observation. We include other common identifiers for each target and a published spectral type, if available. We include our calculated temperature index for the star, which is discussed further in §7. Spectral types and temperature indices of the S stars in the sample may vary as a function of time, because the stars themselves do. When spectral types for different epochs were available, we match the V magnitudes of the epochs to those listed in Table 2 to find the spectral type. Finally, we include identifications of some stars as intrinsic or extrinsic (denoted respectively by 'i' and 'e') from Yang et al. (2006) or from our own calculations based on AKARI magnitudes. For more information regarding the intrinsic/extrinsic identifications, see §6.

2.2. Observations and Data Reduction

Spectra were obtained using the FAST instrument on the 1.5 m Tillinghast reflector on Mt. Hopkins. FAST produces medium-resolution optical wavelength spectra, spanning a wavelength range from 3474 to 7418 Å. All data were taken using the 3" slit and a 300 lines mm^{-1} grating, producing a resolution of $\lambda/\Delta\lambda \sim 1,000$ at 6,000 Å. Exposure times varied widely based on the magnitude of the star, but ranged from ~ 3 to 300 s. The two-dimensional spectra were bias-subtracted and flat-fielded by the observatory standard data reduction procedure. After extraction, and the one-dimensional spectra were wavelength calibrated.

We used standard star spectra taken each night of the observations to produce sensitivity curves and flux calibrated spectra. An extinction correction file from Kitt Peak National Observatory was also applied to correct for major atmospheric

effects. The standard calibration stars used to create the sensitivity curve are BD+26 2606, BD+33 2642, BD+17 4708, Feige 34, and HD 19445. Flux calibration and extinction corrections were performed using the *onedspec* package and the *calibrate* task in IRAF. We used a combination of observatory logs and night-to-night comparison of the sensitivity function to determine whether the calibrated target spectra were reliable enough to use in color analysis. We rejected some of the initial data sample based on concerns about non-photometric conditions and atmospheric distortion. Other nights are included in the final data for analysis because while the data may not be entirely photometric, it remains reliable enough for color analysis because of the non-wavelength dependent nature of most remaining distortions. We discuss the errors in our derived SDSS colors further in §5.1.

Many of the targets are luminous enough in red wavelengths that longer exposure times cause parts of the spectrum to saturate. The wavelengths from 6000 Å onward are of particular concern. However, shorter exposure times limited the amount of information we could derive from wavelengths below 4500 Å because of low signal-to-noise ratios. To create an S star spectral atlas, when possible we interpolated between two different exposure times: a shorter exposure time that produced accurate, unsaturated spectra at the longer wavelengths, and a longer exposure that resulted in higher signal-to-noise at shorter wavelengths. Generally, we determined by hand which spectra were reliable in a particular wavelength region by comparing multiple exposure times and examining long exposures for nonlinear behavior in areas from 6000 Å onward. For stars with only one available exposure time, we rejected a star as saturated if the counts exceeded 60,000. After interpolating between spectra to eliminate saturated regions and rejecting some data, we were left with 46 S star spectra.

3. CREATING THE SPECTRAL ATLAS

We use S star line identifications (Wyckoff & Clegg 1978) to identify major spectral features and select objects with unique characteristics to include in the final spectral atlas, which includes 14 objects marked in Table 1. The objects were chosen to be a representative sample of the larger FAST sample, which includes many fundamentally similar objects. Two sample spectra with associated molecular line identifications are plotted in Figure 1. While the spectra display the same molecular bands, most notably those associated with ZrO, they have different overall colors, which may be due to differences in abundances and/or effective temperatures. We also note that these spectra show some TiO bands, indicating that they are not pure S stars, but instead probably belong to the MS classification. The appearance of TiO bands indicates the possibility of a C/O ratio slightly less than 0.95. The creation of a spectral atlas allows for greater investigation into the similarities and differences within the S star class. For instance, some members of the spectral atlas exhibit, in addition to prominent ZrO features, H α , H β , and H γ emission lines. The presence of these emission lines indicates that these objects are almost certainly AGB stars, e.g., Miras. The spectral atlas is made available digitally as 14 individual FITS files via this journal.

4. SPECTRAL INDICES OF S STARS

We can immediately use the first digital spectral atlas of S stars to derive some medium-resolution spectral indices, potentially useful to classify stars as M, S, or C even where they

may have very similar broadband colors. Apart from our own FAST spectra of 48 S stars, we use M giants from the Indo-US spectral atlas (Valdes F., et al. 2004). We use eight carbon star spectra also obtained with the FAST for a separate project (Green et al. 2011, in preparation). These latter spectra span a wide range of spectral band strengths, but are selected from the SDSS color wedge defined by Margon, B., et al. (2002).

We defined three indices, one each for TiO, ZrO, and C₂. We follow the basic premise of using ratios of mean flux per Ångström across key molecular bands, so that spectra of differing resolution should yield similar results. We use a neighboring comparison region for each band, so that sensitivity to broadband flux calibration is minimized. For the ZrO index, we divide the mean flux Å across 6400 – 6460 Å (ZrO off-band) by that between 6475 – 6535 Å (ZrO band). For the TiO index, we use mean Å⁻¹ flux from 7065 – 7175 Å (TiO band) divided by 6965 – 7028 Å (TiO off-band). The C index is derived from 5025 – 5175 Å (C₂ band) and 5238 – 5390 Å (off-band). Thus, lower flux in the denominator caused by strong molecular bands yields a larger spectral index value. We note that the C index near 5100 Å can also be substantially contaminated by TiO, so we label it “CT51” hereafter. At least within this representative sample of stars, a combination of these indices yields a preliminary S star classification, as shown in Figure 2. The objects most reliably classified as S stars would have ZrO index > 1.2 and C₂ index < 1.4.

5. COLOR SELECTION OF S STARS

5.1. Generating SDSS and 2MASS Colors

To use the SDSS database to find likely candidates for both S giants and dwarfs, we need reliable colors for our objects in as many SDSS bandpasses as possible. The spectra from FAST cover the bandpasses for the *g* and *r* filters completely, and we are thus able to convolve our spectra with these bandpasses and generate SDSS colors. Each of the FAST spectra contains 2681 data points, with approximate wavelength separation of 1.47 Å. To convolve the spectra with the filter transmission curves, we use linear interpolation between transmission values to match a transmission constant to each data point in the spectrum. Once we find a transmission constant for each wavelength, we normalize the transmission curve and then convolve the spectrum with the transmission curve. We use the IRAF task *calcphot* in the *synphot* package of *stsdas* to convolve the spectra with the bandpasses and obtain SDSS colors from our spectra. The SDSS bandpass curves are taken from information made available by the SDSS consortium (Fukugita et al. 1996). As with all SDSS magnitudes, the resultant magnitudes are in the AB magnitude system. By including both SDSS and 2MASS bandpasses, we are able to calculate 10 unique colors.

We also conduct error analysis on the derived SDSS colors using the standard star spectra and the known *g-r* colors and *r* magnitudes for BD+26 2606, BD+33 2642, BD+28 4211, and BD+174708. These colors are published by the SDSS consortium, and allow us to constrain our errors in both the magnitudes and colors (Smith et al. 2002). We calculate the difference between measured and standard *g* and *r* magnitudes, and find a standard deviation of $\sigma_g = 0.191$ mag for the *g* band and $\sigma_r = 0.192$ mag for the *r* band. While these seem high and would suggest, via propagation of errors, a correspondingly high error in our colors (on the order of 0.4 mag), we find a standard deviation of only $\sigma_{color} = 0.029$ mag for

the color. The value was calculated by comparing the derived $g-r$ colors for the standard stars to the known $g-r$ colors from the SDSS Data Release 5 standard star network (Smith et al. 2002). The low value in the uncertainty of our colors is probably due to the fact that while our spectra may not be absolutely photometrically calibrated, many of the remaining distortions are not wavelength dependent and therefore cancel out during color calculation. Examples of remaining distortions include slit losses, an incomplete cancellation of the telluric spectrum or a thin cloud layer that dims the overall flux at all wavelengths. So while individual magnitudes remain unreliable at best, the errors in the derived color are comparable to color errors in 2MASS and are thus reliable enough for our analysis. The 2MASS magnitude errors for our sample have a wide range dependent upon the reliability of the photometry for a particular night, but generally range from 0.02 magnitudes up to about 0.2 magnitudes. Colors calculated from 2MASS magnitudes thus have errors comparable to the derived SDSS colors.

5.2. Color Selection Analysis and Criteria

The initial aim of the project was to select for S stars based on their SDSS and 2MASS colors. To achieve this goal, we derive g and r magnitudes for the stars making up our sample and combine with the known 2MASS magnitudes to construct a total of 10 unique colors. The g , r , and 2MASS magnitudes, along with V band, IRAS, and AKARI photometry where available, are presented in Table 2. Many of these colors, however, are ultimately unhelpful for analysis purposes because the colors of known S stars cover several magnitudes. Such a wide spread in magnitudes does not allow for efficient discovery of the S stars based on color. In other color combinations, the S stars are not well distinguished from the stellar locus, so selection based on color is contaminated by large numbers of stars from the main sequence. We find that the most useful color combinations are $g-r$, $J-H$, and $H-K_s$, because in these colors, the FAST S stars sample is well distinguished from the stellar locus. We compare these colors with those found for general SDSS stars, by overplotting some 300,000 SDSS/2MASS matches of (Covey et al. 2007), which they used to parametrize the stellar locus in terms of color. In Figure 3, we present two color-color diagrams in which the S giants from the FAST sample are best separated from the stellar locus. The first shows the comparison of $g-r$ color and $J-H$ color, while the second shows the comparison $r-H$ and $J-K_s$ color.

However, the S giants that comprise the FAST sample, by virtue of their location on the TP-AGB or RGB, have similar colors and effective temperatures to M giants. In Figure 3, we also examine colors for M giants and dwarfs matched in the SDSS and 2MASS surveys. While the M giants are very few in number (due to the depth of both of these surveys, many nearby M giant targets are saturated), we can immediately see that S stars and M stars lie in essentially the same area of color space. We also include C stars on the color-color diagrams, since they follow evolutionary paths similar to S stars and also have enhanced abundances. The S stars of the FAST sample, while well-distinguished from the overall stellar locus on both color-color diagrams, are particularly difficult to distinguish from C stars on the $g-r$ versus $J-H$ diagram. On the plot comparing $r-H$ to $J-K_s$, S stars can be distinguished reasonably well from the bulk of the M star population using $J-K_s > 1.2$ and $r-H > 3.5$. Given the biases in the current sample of S stars, it is difficult to determine the efficiency of

this selection, but this region on the color-color plot appears to be the least contaminated by M and C stars. Spectroscopy for a sample of perhaps 100 stars in this color region could directly test the efficiency of this color selection.

6. DISTINGUISHING INTRINSIC AND EXTRINSIC S STARS

6.1. Techniques

Distinguishing intrinsic and extrinsic S stars is difficult without high-resolution spectroscopy, where direct detection of Tc is possible. However, as presented in Van Eck and Jorissen (2000), there are other measurements which serve to segregate the two types with reasonable statistical accuracy. Since extrinsic S stars are generally on the RGB (or early AGB), they cluster closer to the main sequence than intrinsic S stars in terms of color. Stars on the AGB generally show a red excess which is absent on the RGB. However, Van Eck and Jorissen (2000) find that blue excursions of Mira variables, which make up a significant proportion of intrinsic S stars, make the two classes difficult to distinguish based on distance from the main sequence alone. Additionally, intrinsic and extrinsic S stars have been shown to segregate well in a ($K-[12]$, $K-[25]$) diagram (Groenewegen 1993; Jorissen et al. 1993). These colors are based on IRAS photometry, which is available for many of the stars in the FAST sample. IRAS magnitudes in these bands are presented in Table 2 for 37 stars (with two bands available for 29). Identifications of many of the S stars in the FAST sample as intrinsic or extrinsic are included in Table 1. All such identifications not marked with an asterisk are from Yang et al. (2006). Classification of most of the S stars in our sample is inconclusive because of their intermediate positions in the color-color diagrams. However, this IR technique remains the most efficient photometric classification method known to determine whether S stars are intrinsic or extrinsic.

We also use results from Vanture et al. (private communication), which indicate that S stars presenting Li in their spectra are intrinsic S stars. Lithium in stars is destroyed at fairly low temperatures (Bodenheimer 1965) but can be produced by the Cameron-Fowler mechanism, in high luminosity ($M_{\text{bol}} < -6$) TP-AGB stars via hot bottom burning. Since, as mentioned above, extrinsic S stars are often less evolved along the giant branch than intrinsic S stars, they generally lack any evidence for the resonance line Li 6707Å (Vanture et al. 2011, in preparation). We compare the list of S stars with and without Li detections to our FAST sample and analyze them for the Li 6707Å feature. Unfortunately, we find that this line cannot be reliably detected with spectra of our resolution and S/N.

6.2. Identification Using AKARI Magnitudes

The AKARI satellite surveyed much of the sky in both near and far infrared bands. We use flux data collected by AKARI in the S9W and L18W bandpasses to find infrared magnitudes for six objects in the FAST sample without available IRAS colors (Murakami et al. 2007; Ishihara et al. 2010). The S9W and L18W bandpasses are qualitatively similar to the IRAS 12 and 25 micron bandpasses. All available AKARI magnitudes are presented in Table 2, for 52 stars (with two bands available for 35). Since the zero magnitude fluxes for these bandpasses are not well established, all magnitudes are presented on the AB system. In Figure 4, we show two color-color diagrams comparing intrinsic and extrinsic stars using $K-[9]$, $K-[18]$, and $[9]-[18]$ colors. S stars previously classified as intrinsic

vs. extrinsic are shown in black, with others plotted in red.

From the AKARI magnitudes and plots, we are able to identify an additional two stars in the sample as intrinsic or extrinsic. We identify 22315839+0201206 as an extrinsic S star. This star has a $K - [18]$ value of -5.92 mag and a $[K] - [9]$ value of -4.23 mag. This places it in the extreme lower left corner of the color-color diagram comparing $K - [9]$ color to $K - [18]$ color, far away from any known intrinsic stars but close to many known extrinsic stars. Similarly, we are able to identify 19364937+5011597 as an intrinsic S star on the basis of color. This star has a $K - [9]$ value of only -3.09 mag, and a $K - [18]$ value of -4.01 mag. This places it in the upper right-hand corner of the plot comparing $K - [9]$ and $K - [18]$ color, far away from the cluster of extrinsic S stars. The four other stars for which we have no classification (03505704+0654325, 10505517+0429583, 13211873+4359136, and 16370314+0722207) cluster close to the transition between intrinsic and extrinsic S stars in all three color comparisons, meaning that we cannot confidently identify these with either class of stars. We include our additional identifications in Table 1, marked with an asterisk to distinguish them from those from Yang et al. (2006).

7. CALCULATING TEMPERATURE INDICES

7.1. Using Colors to Determine Temperature

As mentioned before, S stars cover a wide range of spectral types and effective temperatures. However, they generally have effective temperatures similar to those of M giants. The starting point for classifying the spectral type and temperatures of our S stars, therefore, is to calculate a temperature index using M-giant criteria (Gray and Corvally 2009). Houdashelt et al. (2000) use a grid of stellar models to calibrate a relation between color and temperature index for M stars. Their analysis uses CIT/CTIO colors $V - K$ and $J - K$. We find V magnitudes for some of the FAST sample using Simbad. We also convert the 2MASS K_s magnitudes and $J - K_s$ colors using the following relations from Carpenter (2001):⁴

$$\begin{aligned} K_s &= K_{CIT} + (-0.019 \pm 0.004) + (0.001 \pm 0.005)(J - K)_{CIT} \\ J - K_s &= (1.068 \pm 0.009)(J - K)_{CIT} + (-0.020 \pm 0.007) \end{aligned} \quad (1)$$

We then use the temperature index/color relation to assign a preliminary temperature index to the stars in the FAST sample. Where possible, we use the $(V - K)_{CIT}$ relation, since this is specified by the authors as the most temperature sensitive. When there is no V magnitude available, we use the $(J - K)_{CIT}$ relation (Houdashelt et al. 2000).

7.2. Results of Temperature Analysis

Results of this analysis are presented in Table 1. These temperature indices should be treated with caution, since many of the stars in the FAST sample are highly variable. Furthermore, differences in C/O ratios, may lead to temperature errors up to 400 K (Van Eck et al. 2010). However, the preliminary classification gives us a rough idea of the relative effective temperatures of the S stars. We find that the average temperature index of the stars in the sample is 5, which in M giants corresponds to an effective temperature of roughly 3500 K (Houdashelt et al. 2000). In Table 1, we also present known spectral types from Keenan & Boeshaar (1980), which

include temperature indices for the S stars. We find reasonably good agreement between our calculated temperature indices and those presented as part of the spectral type. Differences between our temperature index and those included as part of the published spectral types are within the range of variability of the star. We also note that since we do not have a reliable color/temperature relation for stars with a temperature index greater than 7, we present all of these classes as 7+. We also find reasonably good agreement (within one temperature index) between the values using $(V - K)_{CIT}$ and those derived using $(J - K)_{CIT}$.

We also conduct a preliminary analysis on possible correlations between the intrinsic/extrinsic distinction and the temperature indices. We find that the average temperature index of the intrinsic stars is 5.6 ± 2.3 , while the average temperature index of extrinsic stars is 4.3 ± 2.2 . Therefore, while the mean temperature index shows some differences, this measurement alone is not statistically significant. A larger sample would be needed to determine the validity of differences between average effective temperature between intrinsic and extrinsic S stars. The temperature indices of S stars in general are insufficient to classify them as intrinsic or extrinsic, except for the reddest objects.

8. PARALLAX, PROPER MOTION, AND ABSOLUTE MAGNITUDE

8.1. Parallax Analysis

Four of the stars in the original FAST sample have reliable parallax measurements (defined as a parallax detection at the 3σ level) available from the Hipparcos survey (Perryman et al. 1997), so calculation of the absolute magnitude of these stars in the g and r bands is possible. We calculate the error in our absolute magnitudes using the presented error in the parallax measurements and assuming apparent magnitude error values of $\sigma_g = 0.191$ and $\sigma_r = 0.192$. Results of this analysis are given in Table 3. We note the calculated absolute g and r magnitudes, as well as the associated errors. Objects are identified using the 2MASS Identifier, and the parallax and associated error as given by the Hipparcos catalog are also included. We can immediately identify each of these four objects as asymptotic or red giant branch members: if they were dwarf stars of the same spectral type, we would expect magnitudes in the g and r of $+5$ or higher, as opposed to the derived values, which center around 0.

8.2. Proper Motion Analysis

We also find that some stars in the original FAST sample display significant (again, 3σ) proper motions. The proper motions are available in the Tycho-2 catalogue (Høg et al. 1997). Of these, we have reliable g and r photometry for 16 objects. Given the proper motion and assuming that all objects are moving at total space velocities less than the Galactic escape velocity, we can derive an upper limit on the distance, and thus also constrain the absolute magnitudes of the objects. To derive an approximate distance (d), we use the equation:

$$d = \frac{v_{trans}}{\mu} < \frac{v_{esc}}{\mu} \quad (2)$$

$$d < 0.0056 \text{ pc/yr} \left(\frac{\mu}{1 \text{ rad/yr}} \right)^{-1} \quad (3)$$

$$d < 1.16 \text{ kpc/yr} \left(\frac{\mu}{1 \text{ mas/yr}} \right)^{-1} \quad (4)$$

⁴ Updated at <http://www.astro.caltech.edu/~jmc/2mass/v3/transformations>

where μ is the proper motion of the object measured in radians per year, v_{trans} is the transverse velocity of the object, and v_{esc} is the Galactic escape velocity. We take the Galactic escape velocity to be 545 km/s (Smith et al. 2007), which is equivalent to 0.0056 parsecs per year. The distance constraint is only an upper limit because we neglect radial velocities. However, we note that our *rvsao* measurements generally yield small radial velocities ($\lesssim 50$ km/s) for the S star sample. We then use the same formulas presented above to derive g and r absolute magnitudes and associated errors. These g and r absolute magnitudes essentially represent the brightest the object concerned could be, considering its proper motion and assuming that all observed objects are gravitationally bound members of the Milky Way. Results of the analysis are presented in Table 4. We can immediately see from these lower bound magnitudes that these objects could be RGB or AGB stars - even considering that the actual magnitudes are probably significantly fainter. For the sample of 16 stars, we find that the average lower limit on the absolute g magnitude M_g is -5.0 mag, correspondingly $M_r \sim -6.7$ mag. Cool dwarf stars generally have g and r absolute magnitudes greater than $+8$ mag (Dahn et al. 1977; Vyssotsky 1947). We conclude that all 16 objects displaying significant proper motions could be giants.

8.3. Lower Limit Distances

To further bear out this analysis, we also determine a lower limit on the distance for these objects by assuming (conversely to above) that they *do* have a dwarf magnitude. We take the approximate absolute g -band magnitude for an M dwarf star to be $M_g = 10$ mag and the absolute r -band magnitude to be $M_r = 9$ mag (Dahn et al. 1977; Vyssotsky 1947). This method can be applied to the entire sample, because it does not rely on proper motion or parallax measurements. We find that most of these objects, were they to have typical dwarf magnitudes, would be within 25 pc of us. This distance would almost guarantee that these objects would have detectable parallax and proper motion. For instance, assuming extrinsic S stars are a mostly spheroid population (like dC or CH stars: see Green et al. (1994) and Bergeat et al. (2002)) with velocity of -220 km/s relative to the Sun (Sirko et al. 2004), the typical proper motion for a star 25 pc away would be on the order of 2.25 mas/yr. Additionally, such stars would have a parallax on the order of 40 mas. Since such a parallax would be easily detectable by Hipparcos, these lower limit distances therefore also suggest that the stars of the FAST sample are likely to be giants or AGB stars.

9. SUMMARY AND CONCLUSIONS

Using a sample of 57 medium-resolution S star spectra taken with the FAST spectrograph, we create a spectral atlas of S stars comprised of 14 objects that span a range of spectral types within the MS, S, and SC classes. The atlas is published as a collection of 1-D FITS files via this journal. After generating g and r SDSS magnitudes from the spectra, we find that the SDSS magnitudes in the g and r bands, combined with J , H , and K_s magnitudes from the 2MASS catalog for S stars may with further confirmation allow for reasonably efficient color selection of these objects from the SDSS and 2MASS catalogs. We use previously published data to identify some of the stars in the sample as intrinsic or extrinsic stars, and find that the fraction of extrinsic S stars in the sample is approximately 54%. We also assign temperature indices to the stars in the sample based on the M star scale of temperature indices. We find that much of the S star sample falls at temperature index 4 or above, meaning that the effective temperatures of most of these stars are well below 5000 K. We also analyze objects in the FAST sample with detectable parallaxes and proper motions to generate absolute magnitude limits, as well as lower limit distances based on assumed dwarf magnitudes. This analysis bears out our initial assumption that the FAST sample is primarily composed of giant stars, either on the AGB or RGB.

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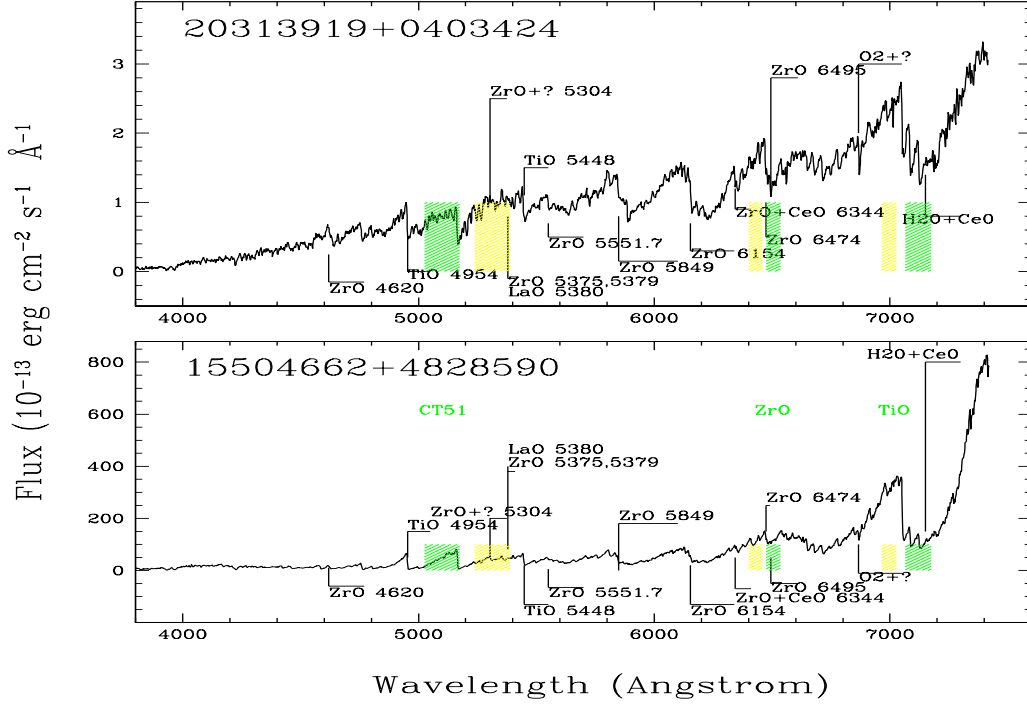


FIG. 1.— FAST spectra of two different S stars, which are named by their 2MASS identifiers. The bandheads of major absorption features are labelled. Flux is in units of $10^{-13} \text{ ergs s}^{-1} \text{ cm}^{-2} \text{ Å}^{-1}$. While there are few bands unique to one of the spectra, the colors of two objects are quite different. The wavelength ranges we use for spectral indices are shown as colored bars, green within the molecular absorption band, yellow for the nearby comparison bandpass. Labeled in green near 7000Å are the TiO index bandpasses, near 6500Å are the ZrO index bandpasses, and near 5200Å are the bandpasses for the “CT51” index, which is intended to measure C_2 but may contain substantial effects from TiO.

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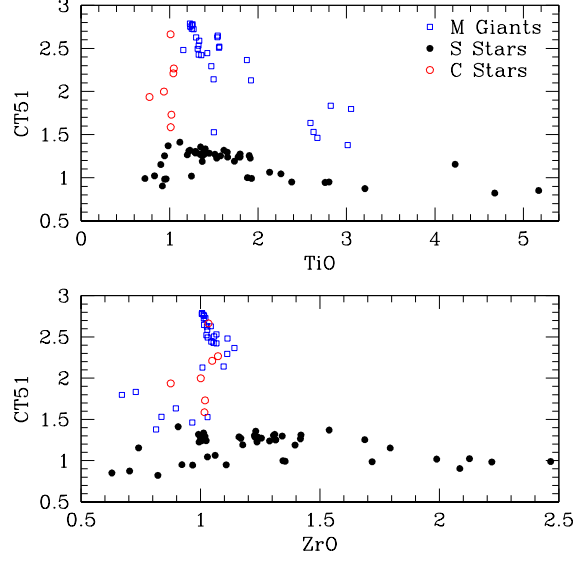


FIG. 2.— Spectral indices of S, M, and C stars. (*top*) CT51 vs. TiO indices for M giants (open blue boxes), S stars (filled black circles) and C stars (open red circles). M and C stars show ZrO indices below ~ 1.2 . S stars show a wide range of ZrO indices, but have CT51 indices less than ~ 1.4 .

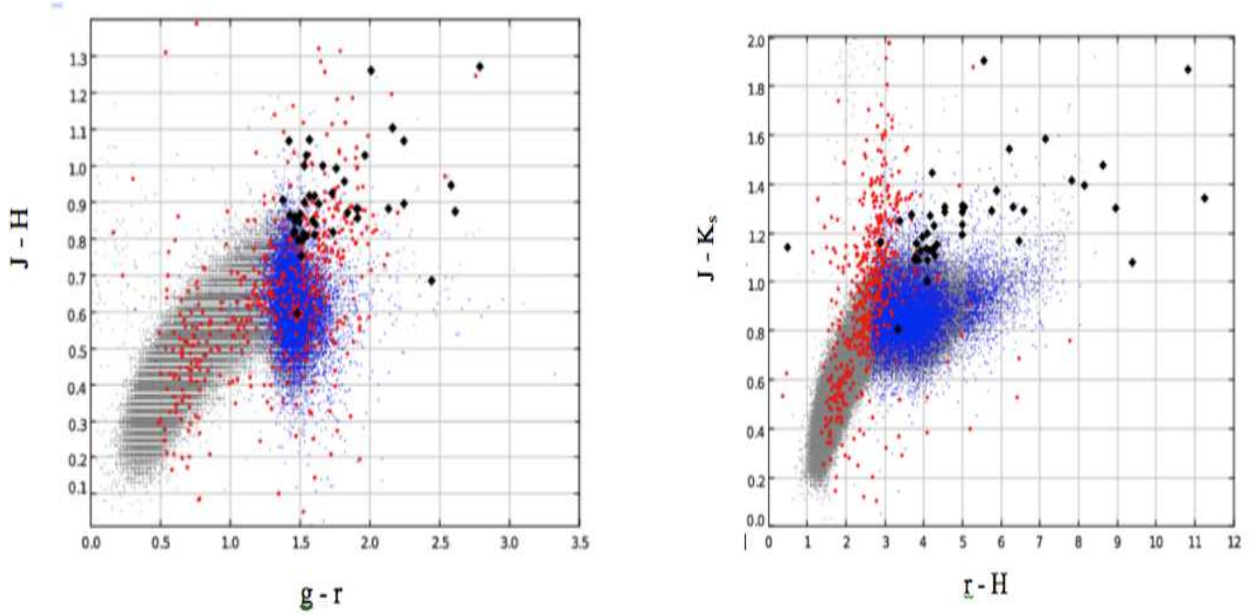


FIG. 3.— Two color-color diagrams, comparing the stellar locus (small grey points), the FAST S star sample (black), C stars (red), and known M stars (in blue). Carbon star colors are from SDSS and 2MASS matches, as presented by Green et al. (in preparation). (*left*) A plot of $g-r$ color versus $J-H$ color. (*right*) A plot of $r-H$ versus $J-K_s$ color.

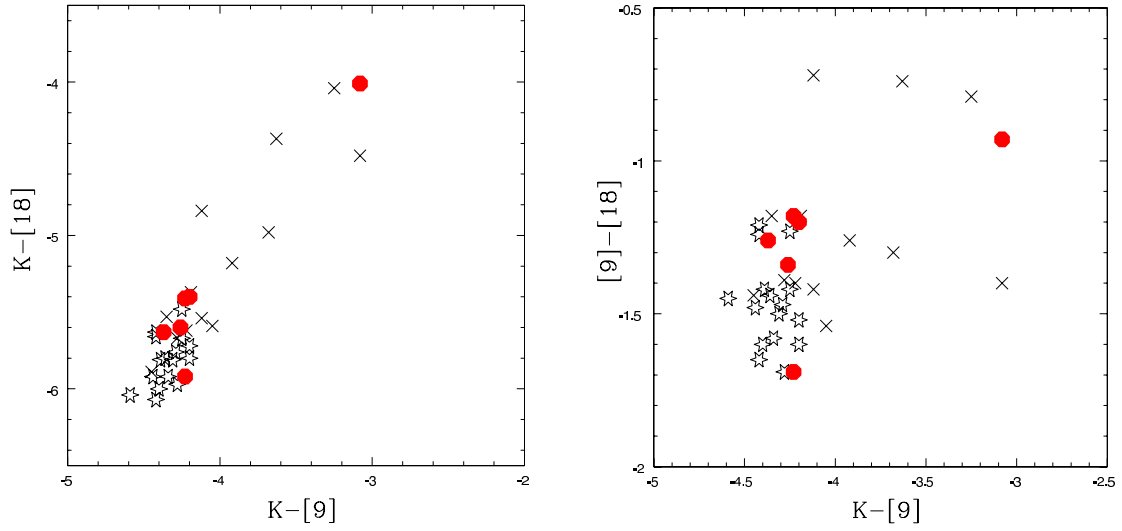


FIG. 4.— A comparison of 2MASS K magnitudes and AKARI infrared magnitudes to further distinguish intrinsic and extrinsic S stars. Known intrinsic stars are plotted as crosses, while known extrinsic stars are plotted with star symbols. Previously unclassified objects are plotted in red. These plots show distinctions between intrinsic and extrinsic S stars similar to those in Figure 4, which uses IRAS magnitudes. (*top left*) A plot comparing $K-[9]$ and $K-[18]$ color. (*top right*) A plot comparing $K-[9]$ and $[9]-[18]$ color.

TABLE 1
TARGET S STARS

Stephenson Catalog Number	RA J2000 (deg)	Dec J2000 (deg)	2MASS Identifier	Other Identifiers	Obs. Date	Spectral Type	Temp. Class	Intrinsic vs. Extrinsic
9	006.008240	+38.577049	00240197+3834373 ^A	HD 1967/R And	09/15/09	S5e Zr5 Ti2	7	i
22	016.300978	+19.197842	01051223+1911522	HD 6409/CR Psc	09/15/09		4/5	e
32	021.413226	+21.396074	01253917+2123458	RX Psc	09/15/09		0/1	i
45	028.582236	+21.889090	01541973+2153207	BD+21 255	09/15/09	S3 Zr1 Ti3	3/4	e
57	036.476487	+38.122776	02255435+3807219	BI And	09/15/09	S8 Zr7 Ti4	7	i
73	052.040388	+17.679628	03280969+1740466	...	09/20/09		4/5	...
74	052.936921	+04.695488	03314486+0441437 ^A	...	09/20/09		6	...
80	055.889249	+22.437023	03433341+2226132	BD+21 509	09/16/09		5/6	e
83	057.737708	+06.909048	03505704+0654325	...	09/16/09		6	...
94	066.090802	-02.532759	04242179-0231579	BD-02 891	09/20/09	S2 Zr2- Ti2	4	e
106	073.937675	+79.999931	04554504+7959597	BD+79 156	09/20/09	S4 Zr2- Ti3-	4/5	e
134	080.836187	-04.570626	05232068-0434142	HD 35273/V535 Ori	09/20/09		7+	i
312	108.717017	+68.804321	07145208+6848155 ^A	HD 54587/AA Cam	11/10/09	M5S†	4/5	i
339	111.489470	+62.591026	07255747+6235276 ^A	...	11/10/09		4	...
347	112.048398	+45.990597	07281161+4559261	HD 58521/Y Lyn	11/10/09	M6S†	7+	i
403	117.325749	+23.734451	07491817+2344040	HD 63334/T Gem	11/10/09	S3e Zr2.5 Ti2	4/5	i
405	117.681704	+47.003967	07504360+4700142	...	11/10/09		4/5	...
413	118.221777	+34.614128	07525322+3436508	BD+34 1698	11/10/09		6	i
418	118.369741	+17.780506	07532873+1746498	HD 64209	11/10/09		4/5	e
431	119.233374	+31.167286	07565600+3110022 ^A	AO Gem	11/10/09		7	i
460	121.836937	+11.321875	08072086+1119187	...	03/10/10		4/5	...
471	122.761565	+08.139297	08110277+0808214	...	03/10/10		5/6	e
494	125.428453	+17.285120	08214282+1717064	HD 70276/V Cnc	03/10/10	S0.5e Zr0	1	i
589	137.661696	+30.963114	09103880+3057472	HD 78712/RS Cnc	03/10/10	M6S†	7+	i
612	143.901826	+69.157043	09353643+6909253	BD+69 524	03/11/10		3/4	e
707	162.729912	+04.499532	10505517+0429583	...	03/11/10		6	...
722	166.970132	+68.366440	11075283+6821591	HD 96360/HL UMa	03/11/10		4/5	e
803	190.986134	+61.093254	12435667+6105357	HD 110813/S UMa	03/11/10	S6e Zr6	5/6	i
819	200.328046	+43.987118	13211873+4359136	AV CVn	03/11/10	S3 Zr2 Ti1	3/4	...
833	207.141756	+31.999107	13483402+3159567	...	03/11/10		0	...
836	208.005660	-03.480082	13520135-0328482	HD 120832	03/11/10		4/5	e
856	214.462975	+83.831696	14175111+8349541 ^A	HD 1272266/R Cam	06/24/09	S2e Zr2	3	e
855	216.950026	-03.079711	14274800-0304469	BD-02 3848	05/03/09		4	e
892	232.384097	+00.189076	15293218+0011206	...	05/03/09		2/3	...
902	238.100360	-23.352097	15522408-2321075	...	05/03/09		6	e
903	237.694254	+48.483078	15504662+4828590 ^A	HD 142143/ST Her	05/03/09	M6.5S Zr1 Ti6+	5/6	i
926	245.401029	+56.877048	16213624+5652373 ^A	HD 147923	05/03/09		2	e
932	249.263106	+07.372439	16370314+0722207 ^A	BD+07 3210	05/03/09		4/5	...
958	256.694247	-02.017769	17064661-0201039	...	05/03/09		6	...
981	260.529640	+23.817570	17220711+2349032	BD+23 3093	05/03/09	S3.5 Zr3+ Ti4	4/5	e
986	261.200788	+33.918953	17244818+3355082	...	05/03/09		3	...
1002	267.252799	+21.710087	17490067+2142363 ^A	...	05/03/09		4	e
1065	279.398924	+43.256588	18373574+4315237	...	05/03/09		3	e
1087	282.875526	+48.911934	18513012+4854429 ^A	TU Dra	05/03/09		3/4	i
1150	294.205739	+50.199917	19364937+5011597	HD 185456/R Cyg	05/03/09	S5e Zr5 Ti0	7	i*
1152	294.345155	+67.188873	19372283+6711199	...	05/03/09	S5- Zr4.5	5/6	i
1165	297.641369	+32.914139	19503392+3254509 ^A	HD 187796/chi Cyg	05/03/09	S6+e Zr2 Ti6.5	7+	i
1222	307.913302	+04.061802	20313919+0403424 ^A	...	05/23/09		5	...
1247	313.355676	+06.539369	20532536+0632217	...	05/23/09		2/3	...
1262	317.601674	+01.606119	21102440+0136220	...	05/27/09		4/5	...
1264	318.232437	+09.548880	21125578+0932559	...	07/17/09		3/4	...
1290	331.995242	+29.665623	22075885+2939562	...	07/17/09		4/5	e
1299	337.993328	+02.022399	22315839+0201206	...	07/17/09		5	e*
1304	340.954936	+33.926231	22434918+3355344	HD 215336	07/17/09		1/2	e
1315	343.648463	+16.941910	22543563+1656308	HR Peg	07/19/09	S4+ Zr1.5 Ti4	4/5	i
1328	349.035110	+28.863674	23160842+2851492 ^A	...	07/19/09		3	e
1340	356.540618	+34.783325	23460974+3446599	...	07/19/09		3	...

NOTE. — The 58 S stars targeted by the FAST survey. The Stephenson Catalog number corresponds to the object's identification in The General Catalog of S Stars, second edition (Stephenson 1984). The 14 stars included in our digital spectral atlas are marked by an A in the 2MASS Identifier. Targets were initially selected from this survey and then matched with 2MASS counterparts (Skrutskie et al. 2006) to provide the J, H, and K_s magnitudes. Spectral types are taken from Keenan & Boeshaar (1980) wherever possible. In some cases, they show types for multiple epochs, in which case we chose the epoch whose listed V mag in Skiff et al. (2010) is closest to the V mag we list in Table 2. If those types are not available, we list types from Keenan (1954) as † and Ake (1979) as ‡. We also present our own calculations of the effective temperature, here presented on the M giant scale, as described in §7.1. Intrinsic/extrinsic identifications denoted with an asterisk are derived using AKARI, rather than IRAS, photometry. For more discussion, see §6.2.

TABLE 2
S STAR MAGNITUDES

2MASS Identifier	<i>g</i>	<i>r</i>	<i>V</i>	<i>J</i>	<i>H</i>	<i>K_s</i>	[S9W] AB	[12]	[L18W] AB	[25]
00240197+3834373	8.40	6.37	7.39	2.02	0.77	0.12	...	-2.66	3.17	-2.3
01051223+1911522	9.38	7.79	...	3.67	2.76	2.48	6.68	1.75	8.28	1.63
01253917+2123458	16.53	14.40	8.80	6.27	5.39	4.98	8.90	3.91	10.16	...
01541973+2153207	10.18	8.60	9.00	5.29	4.48	4.29	8.57	3.59	10.26	3.61
02255435+3807219	11.90	9.63	...	3.50	2.43	1.92	6.14	1.10	7.54	0.95
03280969+1740466	12.81	11.14	...	7.91	7.02	6.72	11.09
03314486+0441437	12.68	11.05	...	6.93	6.02	5.65	9.92
03433341+2226132	10.74	9.05	9.69	5.03	4.03	3.80	8.14	3.12	9.72	3.23
03505704+0654325	12.40	10.54	...	6.42	5.47	5.12	9.32	4.32	10.52	...
04242179-0231579	10.42	8.78	9.33	5.62	4.78	4.44	8.80	3.96	10.24	3.85
04554504+7959597	10.68	9.04	9.66	5.67	4.87	4.55	8.97	4.04	10.18	3.99
05232068-0434142	11.23	9.04	10.00	3.90	2.80	2.36	...	0.38	6.47	-0.50
07145208+6848155	8.85	7.33	...	2.60	1.67	1.39	5.74	0.70	6.92	0.08
07255747+6235276	12.49	10.92	...	7.67	6.83	6.53	10.79
07281161+4559261	7.37	0.65	-0.38	-0.69	3.43	-1.59	4.15	-0.67
07491817+2344040	13.37	11.10	8.00	4.12	3.22	2.71	7.16	2.20	8.60	2.05
07504360+4700142	11.83	10.26	...	7.32	6.43	6.17	10.43
07525322+3436508	11.34	9.79	...	4.99	3.99	3.71	7.99	2.95	9.38	2.81
07532873+1746498	8.56	4.82	3.84	3.47	7.76	2.79	9.23	2.83
07565600+3110022	16.96	14.34	...	6.60	5.66	5.13	8.81	3.73	10.11	2.97
08072086+1119187	10.93	9.43	...	7.37	6.52	6.22	10.43
08110277+0808214	10.19	8.30	...	5.45	4.58	4.17	8.42	3.52	9.84	3.35
08214282+1717064	7.50	5.08	4.00	3.59	6.67	1.46	8.07	1.46
09103880+3057472	6.88	4.94	6.08	-0.71	-1.56	-1.87	...	-3.07	2.72	-1.87
09353643+6909253	10.24	8.78	9.30	5.55	4.69	4.42	8.84	3.84	10.08	3.89
10505517+0429583	12.65	10.90	...	5.36	4.54	4.05	8.28	...	9.46	...
11075283+6821591	9.13	7.73	8.10	4.07	3.17	2.77	7.36	2.37	8.81	2.26
12435667+6105357	9.69	7.69	8.87	4.46	3.43	3.02	7.21	2.06	8.39	1.71
13211873+4359136	11.00	9.37	9.78	6.28	5.43	5.16	9.53	...	10.79	...
13483402+3159567	13.30	11.80	...	9.03	8.43	8.22
13520135-0328482	9.59	5.50	4.76	4.35	8.75	3.78	10.35	...
14274800-0304469	10.26	8.77	9.47	5.74	4.93	4.59	9.01	4.05	10.66	...
14175111+8349541	10.58	8.81	6.97	3.82	2.90	2.46	6.90	1.83	8.38	1.66
15293218+0011206	7.47	6.70	6.40	10.77
15522408-2321075	6.67	5.76	5.38	9.68	4.79
15504662+4828590	8.41	6.50	...	0.74	-0.14	-0.54	3.09	-2.12	3.83	-2.9
16213624+5652373	8.50	7.05	7.59	4.71	3.65	3.46	7.66	2.82	9.18	2.64
16370314+0722207	10.50	8.95	9.53	5.50	4.71	4.38	8.64	...	9.98	...
17064661-0201039	12.73	10.94	...	7.35	6.36	6.07
17220711+2349032	10.91	9.37	10.10	5.73	4.98	4.59	8.98	4.09	10.40	3.99
17244818+3355082	12.28	10.73	...	7.73	6.92	6.63	10.98
17490067+2142363	12.71	11.22	...	7.78	6.94	6.64	10.95	5.98
18373574+4315237	12.42	10.93	...	7.46	6.61	6.36	10.61	5.86
18513012+4854429	17.02	14.59	10.00	5.82	5.14	4.74	7.99	2.52	8.78	2.01
19364937+5011597	12.21	9.58	8.15	2.25	1.38	0.86	3.94	...	4.87	...
19372283+6711199	11.18	9.23	9.70	5.07	4.19	3.76	7.88	2.79	9.30	2.64
19503392+3254509	12.55	9.79	6.80	0.17	-1.10	-1.70	...	-4.44	...	-1.68
20313919+0403424	12.43	10.90	...	7.48	6.61	6.25	10.64
20532536+0632217	12.78	11.29	...	7.99	7.17	6.90	11.14
21102440+0136220	7.60	6.79	6.44	10.71
21125578+0932559	7.94	7.09	6.81	11.15
22075885+2939562	6.56	5.71	5.39	9.64	4.90	10.87	...
22315839+0201206	6.46	5.56	5.21	9.44	...	11.13	...
22434918+3355344	7.82	4.98	4.07	3.79	8.10	3.18	9.60	3.02
22543563+1656308	7.02	5.43	6.47	2.31	1.24	1.04	5.09	-0.02	6.63	-0.08
23160842+2851492	13.62	12.10	...	9.11	8.32	8.03
23460974+3446599	12.02	10.53	...	7.49	6.63	6.40	10.70

NOTE. — Magnitudes of the S stars in the FAST sample. We present our derived SDSS *g* and *r* magnitudes, as well as the 2MASS magnitudes. We include IRAS [12] and [25] and AKARI [S9W] and [L18W] magnitudes when available. Note that the AKARI magnitudes are calculated using the AB system.

TABLE 3
DISTANCES AND ABSOLUTE MAGNITUDES

2MASS Identifier	Parallax (mas)	σ_p (mas)	Distance (pc)	σ_d (pc)	abs. g (mag)	σ_g (mag)	abs. r (mag)	σ_r (mag)
09103880+3057472	8.06	0.98	125	15	1.41	0.33	-0.53	0.33
15504662+4828590	3.07	0.75	326	80	0.85	0.56	-1.06	0.56
16213624+5652373	2.31	0.62	433	117	0.32	0.61	-1.13	0.61
22543563+1656308	3.31	0.93	302	85	0.02	0.64	-1.97	0.64

NOTE. — Parallax, distance, and absolute magnitudes for four members of the FAST sample of S Stars with detectable (3σ) parallax measurements from Hipparcos.

TABLE 4
PROPER MOTIONS, AND DERIVED DISTANCE AND MAGNITUDE LIMITS

2MASS Identifier	Proper Motion (mas/yr)	σ_{pm} (mas/yr)	Distance ^{lim} (kpc)	g mag (mag)	r mag (mag)	M_g^{lim} (mag)	M_r^{lim} (mag)
00240197+3834373	35.89	2.46	3.22	8.4	6.37	-4.14	-6.17
01051223+1911522	12.65	0.95	9.13	9.38	7.79	-5.42	-7.01
03505704+0654325	10.66	2.50	10.8	12.4	10.55	-2.77	-4.62
04242179-0231579	7.50	1.27	15.4	10.41	8.78	-5.53	-7.16
07145208+6848155	18.44	1.07	6.26	8.85	7.33	-5.13	-6.65
09103880+3057472	34.22	0.85	3.38	6.8	4.94	-5.84	-7.70
09353643+6909253	8.54	1.67	13.5	10.24	8.78	-5.42	-6.88
11075283+6821591	22.65	1.15	5.10	9.13	7.73	-4.41	-5.81
12435667+6105357	13.62	1.31	8.48	9.69	7.69	-4.95	-6.95
13211873+4359136	11.31	2.18	10.2	10.99	9.37	-4.06	-5.68
14175111+8349541	7.26	1.62	15.9	10.58	8.82	-5.43	-7.19
14274800-0304469	13.18	1.89	8.77	10.26	8.77	-4.45	-5.94
15504662+4828590	22.38	1.11	5.16	8.41	6.5	-5.15	-7.06
16213624+5652373	12.32	1.14	9.37	8.5	7.05	-6.36	-7.81
19372283+6711199	9.75	2.35	11.8	11.18	9.24	-4.19	-6.13
22543563+1656308	15.75	0.99	7.33	7.02	5.43	-7.31	-8.90

NOTE. — Stars from the FAST sample with significant proper motions and derived quantities. Distances are upper limits, as described in § 8.2, derived by assuming $v_{trans} < 545$ km/s. Absolute magnitudes are the corresponding lower (brightest) limits. We can immediately see that these stars could all be giants. We do not present errors in the distance and magnitude calculations, since these are rough upper limits, rather than rigorous calculations of the actual expected magnitude.